

## LAMOST J202629.80+423652.0 is not a symbiotic star

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**Abstract.** LAMOST J202629.80+423652.0 has been recently classified as a new symbiotic star containing a long-period Mira, surrounded by dust (D-type) and displaying in the optical spectra high ionization emission lines, including the Raman-scattered OVI at 6825 Å. We have observed LAMOST J202629.80+423652.0 photometrically in the BVRI bands and spectroscopically over the 3500-8000 Å range. We have found it to be a normal G8 IV sub-giant star, deprived of any emission line in its spectrum, and reddened by  $E_{B-V}=0.35$  mag. Combining our photometry with data from all-sky patrol surveys, we find LAMOST J202629.80+423652.0 to be non variable, so not pulsating as a Mira. We have compiled from existing sources its spectral energy distribution, extending well into the mid-Infrared, and found it completely dominated by the G8 IV photospheric stellar emission, without any sign of circumstellar dust. We therefore conclude that LAMOST J202629.80+423652.0 is not a symbiotic star, nor it is pulsating or been enshrouded in dust.

**Key words:** binaries: symbiotic – surveys: LAMOST

### 1. Introduction

Symbiotic stars (SySt) are binary systems composed by a red giant star (RG) that transfers material to a degenerate companion which can be either a white dwarf (WD) or a neutron star (NS). SySt are broadly divided into two major groups (see the recent review by Munari 2019): those *accreting-only* whose optical spectra are dominated by the RG with no or weak emission lines, and the *burning-type* displaying a strong nebular continuum and a rich emission line spectrum: they originate in the wind of the RG largely ionized by the very hot and luminous WD companion that is nuclearly burning at the surface in stable conditions. The accreting-only SySt usually requires satellite UV/X-ray observations to be discovered (cf. the prototype SU Lyn, Mukai et al. 2016), while the burning StSy can easily be discovered through the whole Galaxy and the Local Group thanks to the prominent emission lines. In about 20% of the known SySt (see recent catalogs of symbiotic stars by Belczyński et al. 2000, and Akras et al. 2019), the cool giant pulsates as a Mira and this usually comes also with the

presence of warm circumstellar dust ( $\sim 500\text{--}1000$  K), leading to a classification as D-type according to the scheme introduced by Allen (1982).

Li et al. (2015) have announced the discovery of two new SySt during the LAMOST all-sky spectroscopic survey. LAMOST is a 4m Schmidt telescope located in Xinglong (China), feeding light to 16 spectrographs via 4000 optical fibers. The spectrum presented by Li et al. for one of these two new SySt, namely LAMOST J202629.80+423652.0 (hereafter LM), looks weird and we suspected to be the result of problematic sky subtraction, so decided to investigate the matter by acquiring new spectra and supplementing them with acquisition of BVRI photometry. We report in this paper the results of these observations.

## 2. Observations

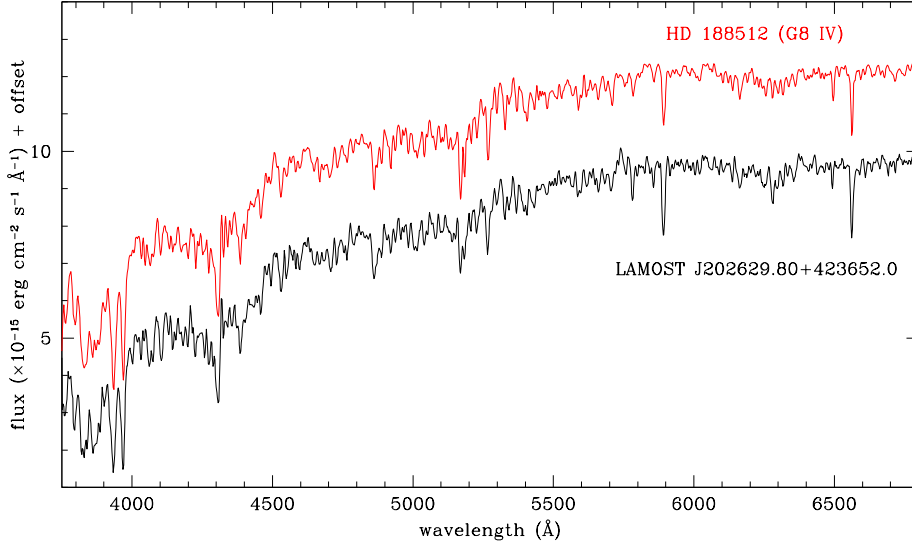
### 2.1. Spectroscopy

Low resolution spectra of LM have been obtained with the 1.22m telescope operated in Asiago by the University of Padova. A Boller & Chivens spectrograph is attached to the Cassegrain focus feeding light to an ANDOR iDus DU440A camera, that houses a back-illuminated E2V 42-100 CCD (2048 X 512 array,  $13.5\ \mu\text{m}$  pixel). A  $300\ \text{ln mm}^{-1}$  grating blazed at  $5000\ \text{\AA}$  was adopted, allowing to cover the  $3300\text{--}8000\ \text{\AA}$  range at  $2.31\ \text{\AA}/\text{pix}$  dispersion. Data reduction was carried out in IRAF and involved the usual steps on bias and dark removal, flat-fielding, variance-weighted spectrum tracing, sky subtraction, wavelength calibration via FeHeAr lamp, heliocentric correction and flux calibration (via observations of the nearby spectrophotometric standard HR 7867). LM has been observed twice on Dec 6 and 14, 2019, exposing for 16min on both occasions. The spectra look identical and their average is plotted in Figure 1. A comparison is there provided with HD 188512, a template for the G8 IV spectral class according to Yamashita et al. (1977), observed with exactly the same telescope setup as for LM, and selected from the Asiago Spectral Database (U. Munari, in preparation).

### 2.2. Photometry

Optical photometry of LM in the Landolt (1992) photometric system has been obtained with a Richey-Chretien 0.4m f/8 telescope located in Monte Baldo (Verona, Italy), equipped with a Moravian G4 CCD camera ( $36\times 36\text{mm}$  active area,  $12\ \mu\text{m}$  size) and B,V,R,I photometric filters from Astrodon. The B filter is of the new version, corrected for the red-leak of its original version (see Munari and Moretti 2012). The photometry of LM that we have collected in five distinct nights is presented in Table 1.

Data reduction has involved all the usual steps for bias, dark, flat with calibration images collected during the same observing nights. We adopted aperture



**Figure 1.** Our spectrum of LAMOST J202629.80+423652.0 compared to that of HD 188512 (+offset), template for the G8IV spectral type.

photometry because the sparse field around LM did not required PSF-fitting procedures. The transformation from the local to the Landolt standard system was carried out via color equations calibrated on a photometric sequence recorded on the same frames as LM:

$$\begin{aligned}
 V &= v + \alpha_v \times (v - i) + \gamma_v \\
 B - V &= \beta_{bv} \times (b - v) + \delta_{bv} \\
 V - R &= \beta_{vr} \times (v - r) + \delta_{vr} \\
 V - I &= \beta_{vi} \times (v - i) + \delta_{vi}
 \end{aligned}$$

The local photometric sequence has been extracted from APASS DR8 survey (Henden and Munari 2014), ported to the Landolt system via the transformations calibrated by Munari et al. (2014). The errors quoted in Table 1 are the quadratic sum of the Poissonian and the transformation contributions.

### 3. Results

#### 3.1. A normal spectral appearance

The spectrum of LM presented in Figure 1 is that of a normal, sub-giant star of the G8IV spectral type, abundant in the Solar Neighborhood as indicated by Gaia DR2 stellar census (Gaia Collaboration 2018). No Mira is known to

**Table 1.** Our BVRI photometry (on the Landolt system of equatorial standards) of LAMOST J202629.80+423652.0.

obs. date	B	V	R	I
2019-12-07.8957		$14.439 \pm 0.009$	$13.766 \pm 0.010$	$13.041 \pm 0.014$
2019-12-14.8155		$14.456 \pm 0.006$	$13.773 \pm 0.007$	$13.119 \pm 0.016$
2020-01-08.8245	$15.629 \pm 0.032$	$14.454 \pm 0.011$	$13.760 \pm 0.010$	$13.100 \pm 0.015$
2020-01-09.8242	$15.603 \pm 0.061$	$14.478 \pm 0.029$	$13.737 \pm 0.017$	$13.076 \pm 0.025$
2020-01-16.7712	$15.580 \pm 0.019$	$14.454 \pm 0.006$		
weighted mean	$15.610 \pm 0.020$	$14.463 \pm 0.004$	$13.755 \pm 0.005$	$13.084 \pm 0.008$

have such an early spectral type (cf. VSX database<sup>1</sup>). In our spectrum of LM no emission line is visible.

Li et al. (2015) report the presence of strong  $H\alpha$ ,  $H\beta$  and [OIII] emission lines in their spectrum of LM. They come actually from the sky-background emission originating from the huge HII region IC 1318 associated to nearby  $\gamma$  Cyg ( $10^\circ$  angular extent). In Figure 2 we present the sky-background at 15arcsec away from LM as recorded on our spectrum. It looks identical to the spectrum attributed to LM by Li et al. (2015).

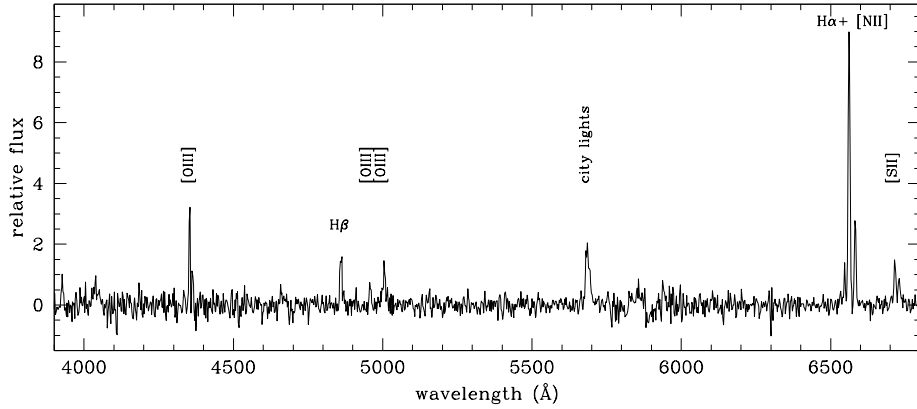
Li et al. (2015) also report about the detection in their spectrum of LM of Raman scattered OVI lines at 6830 and 7088 Å (Schmid 1989). These are actually normal OH lines forming in the Earth's atmosphere, frequently used for wavelength calibration of high resolution spectra at far-red wavelengths (cf. Osterbrock et al. 1996). In Figure 3 we plot the far-red portion of the sky-background at 15arcsec distance from LM taken from our spectrum of the program star, with superimposed the identification of Meinel rotation-vibration bands of OH.

We therefore conclude this section noting that nothing in the spectrum of LM supports a classification as a SySt: all lines and features seen in emission by Li et al. (2015) seems actually due to inaccurate handling of sky-background subtractions.

### 3.2. No photometric variability

Li et al. (2015) also classified LM as a pulsating Mira. Our photometry in Table 1 extends over too short a time interval to draw any conclusion about a possible variability affecting LM. To investigate the matter further, we retrieved the patrol photometry of LM collected by ASASSN (Shappee et al. 2014, Kochanek et al. 2017) and ZTF all-sky patrol surveys (Masci et al. 2019, Bellm et al. 2019). The ASASSN and ZTF photometry in the  $g'$ ,  $r'$ , and  $V$  bands is plotted in Figure 4. It extends from 2015 to 2019. No variability is present in excess of the

<sup>1</sup><https://www.aavso.org/vsx/>



**Figure 2.** Emission lines from the sky-background around LAMOST J202629.80+423652.0 in our spectrum, due to diffuse emission from the HII region around  $\gamma$  Cyg.

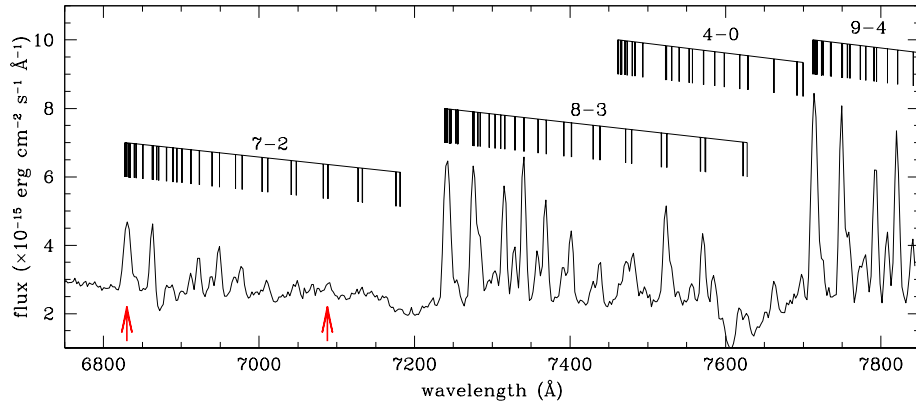
observational noise, and this precludes the classification reported by Li et al. (2015) of LM as a Mira.

### 3.3. No circumstellar dust

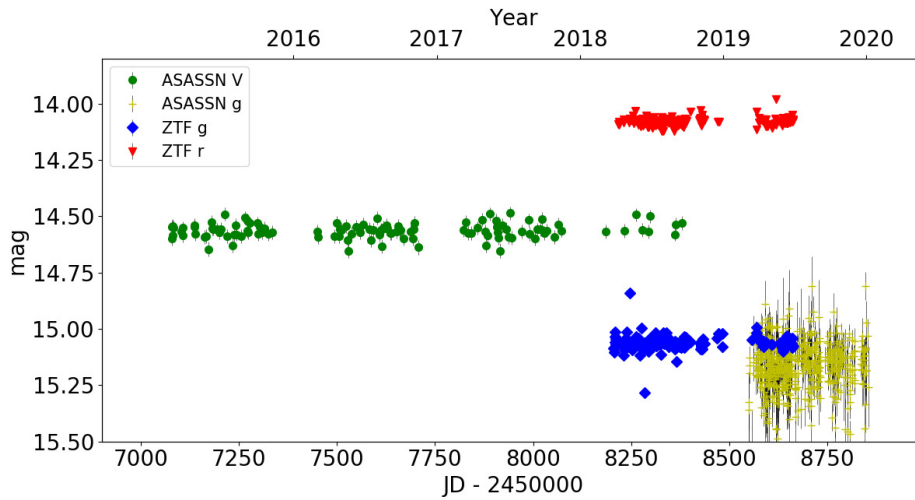
Li et al. (2015) attributed a D-type classification to LM, i.e. the presence of warm dust around it. As illustrated by Allen (1982) the excess emission from such warm dust (500–1000 K) is already evident in the  $H$  band ( $1.6 \mu\text{m}$ ) and become prominent in the  $K$  band ( $2.2 \mu\text{m}$ ). To investigate the spectral energy distribution (SED) of LM, we take advantage of the absence of photometric variability (established in the previous section) that allows to combine non-simultaneous data gathered by different surveys over the whole wavelength range. The SED for LM built from our BVRI observations (Table 1), 2MASS JHK<sub>s</sub> and AllWISE W<sub>1</sub>W<sub>2</sub>W<sub>3</sub> data is presented in Figure 5, together with the reference SED for a typical G8 IV field star using colors from Fitzgerald (1970), Cousins (1976), and Koornneef et al. (1983). The two become a fine match once a correction for  $E_{B-V}=0.35$  reddening is applied to LM. In conclusion, the SEDs in Figure 5 excludes the presence in LM of any dust hotter than  $\sim 100$  K (constrain set by W<sub>3</sub>).

### 3.4. Fundamental parameters

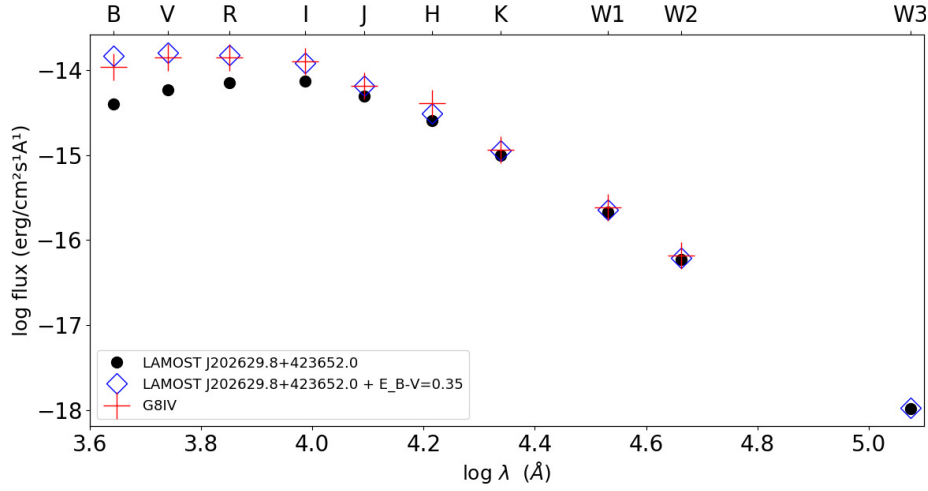
The GAIA second release (GAIA-DR2) lists a parallax  $\pi = 0.761 \pm 0.018$  (mas) for LM. The fractional error is less than 20% and therefore is straightforward to derive the distance to LM as  $d=1.31$  kpc by direct inversion of the Gaia parallax. At such a distance along the line of sight to LM, the Bayestar 3D Galactic



**Figure 3.** Emission in the far-red of the sky-background around LAMOST J202629.80+423652.0 in our spectrum, highlighting the Earth's OH emission (with Meinel rotation-vibration bands identified). The red arrows points to telluric OH lines labelled by Li et al. (2015) as due to Raman scattered OVI 6825, 7088 Å.



**Figure 4.** Photometric data of LAMOST J202629.80+423652.0 obtained by ASAS-SN and ZTF sky patrol surveys.



**Figure 5.** Spectral energy distribution of LAMOST J202629.80+423652.0 as observed (black dots), and after correction for  $E_{B-V}=0.35$  reddening (blue diamonds). The red crosses illustrate the intrinsic energy distribution of a G8 IV star from tabular data.

extinction map (Green et al. 2019) lists a reddening of  $E(g-r) = 0.41$ , that transforms to  $E(B-V)_0 = 0.35$  for a standard  $R_V = 3.1$  interstellar reddening law following the relations for G-type stars by Fiorucci and Munari (2003). This is in excellent agreement with the value of the reddening necessary to bring the SED of LM in agreement with that of a typical G8 IV star in Figure 5. With such a reddening and distance, the mean  $V$ -mag value in Table 1 transforms to an absolute magnitude for LM of  $M(V)=2.78$ . This is in fine agreement with the absolute magnitudes reported by Sowell et al. (2007) that list  $M(V)=5.50$ , 3.20 and 1.35 for G8 stars of luminosity class V, IV and III, respectively. According to tabular values in Straižys (1992), such a star is characterized by a temperature of 5100 K, a luminosity of  $4.4 L_\odot$  and a radius of  $2.7 R_\odot$ .

## 4. Conclusions

We have obtained new photometric and spectroscopic observations of LAMOST J202629.80+423652.0 that combined with literature data shows it to be a normal G8 IV sub-giant star, and not a symbiotic star. Such a classification by Li et al. (2015) appears due to inaccurate sky-background subtraction to their spectra.

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